

Optical spectroscopy and photometry of the neutron star RX J1856.5–3754[★]

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Abstract. We present spectroscopy and imaging with the Very Large Telescope (VLT) of the neutron star RX J1856.5–3754. Little is known about the nature of this source other than that it is a nearby hot neutron star. Our VLT spectrum does not show any strong emission or absorption features. With considerable care to photometric calibration, we obtain photometric measurements over the optical and ultra-violet (UV) using our VLT observations and a detailed analysis of archival *Hubble Space Telescope* data. We find that the entire optical to UV spectral energy distribution is well described by a slightly reddened Rayleigh-Jeans tail ($f_\lambda = (2.96 \pm 0.06) \times 10^{-19} (\lambda/5000 \text{ \AA})^{-4} 10^{-0.4(0.12 \pm 0.05)(A_\lambda/A_V - 1.138)} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$, where A_λ/A_V is the reddening curve; implied $V = 25.58 \pm 0.02$). The reddening is consistent with the interstellar absorption inferred from X-ray spectroscopy. The simplest explanation for this Rayleigh-Jeans emission is that the optical-UV radiation arises from thermal emission from the surface of the neutron star. The high degree to which the data conform to the Rayleigh-Jeans tail significantly limits contributions from other sources of emission. In particular, our observations are inconsistent with the presence of an accretion disk and also strongly constrain the amount of magnetospheric emission from this enigmatic neutron star.

Key words. stars: individual (RX J1856.5–3754) — stars: neutron — X-rays: stars

1. Introduction

The soft X-ray source RX J1856.5–3754 is the brightest and nearest of the so-called isolated neutron stars¹ (for a review, see Treves et al. 2000). These objects have X-ray spectra which appear to be entirely thermal, indicating that the emission arises from the photosphere and that there is little if any contamination from ill-understood emission processes such as those occurring in magnetospheres in radio pulsars and accretion flows in X-ray binaries. Therefore, these objects offer perhaps the best hope of modeling neutron-star spectra, and inferring the effective temperature, surface gravity, and gravitational redshift. In principle, this could lead to unique constraints on the equation of state of matter in the neutron star interiors (e.g., Lattimer & Prakash 2001).

Given these possible gains, ever since its discovery in 1996 by Walter et al., RX J1856.5–3754 has been the sub-

ject of much observational attention. Walter & Matthews (1997) used the *Hubble Space Telescope* (*HST*) to discover a very faint, $V \simeq 25.6$, optical counterpart; its flux is roughly consistent with seeing the Rayleigh-Jeans tail of the $kT \simeq 50 \text{ eV}$ spectrum. Further *HST* observations were used by Walter (2001) to measure the parallax, while Pons et al. (2001) used *HST*, *ROSAT*, *ASCA*, and *EUVE* to measure the broad-band spectral energy distribution.

Pons et al. (2001) also presented detailed model atmospheres for a variety of compositions, with which they were able to model the broad-band spectrum satisfactorily. This leads to strong constraints on the temperature. When combined with the parallax, however, the inferred radii are too small for realistic neutron star models. Pons et al. suggest the surface may not have a uniform temperature distribution. If so, the broad-band spectrum can be used to set only weak constraints on the equation of state.

If one could observe spectral features in the spectrum, one might be able to measure the surface gravity and gravitational redshift without much ambiguity (Paerels 1997). In this respect, the first spectrum at good resolution, taken with *XMM* of RX J0720.4–3125, was disappointing, as no features were seen (Paerels et al. 2001). Recent results

[★] Based on observations collected at the European Southern Observatory, Paranal, Chile (ESO Programmes 63.H-0416 and 65.H-0643).

¹ Given that most radio pulsars are also isolated, a better name might be ‘thermally emitting’ neutron stars.

Table 1. Log of the VLT spectroscopic observations.

Object ^a	UT ^b day, time	Slit or Filter	t_{int} (s)	Par. ^c (°)	sec z
EG 274	15, 23 02	5''	3×10^d	−76	1.21
X-F/+84°6	23 58	B	300	−94	1.67
	16, 00 08	1''	2700	−93	1.60
	00 58	B	3×300	−85	1.33
	01 17	1''	2×2700	−82	1.26
	02 54	5''	300	−53	1.06
X-L/−8°5	03 11	1''	2×2700	−44	1.05
	04 45	5''	300	+37	1.04
X-F/−95°4	05 00	1''	... ^e		
X-F/−95°4	06 06	1''	3×2100	+72	1.15
	07 56	5''	300	+93	1.57
BPM 16274	08 08	5''	3×20^d	−47	1.21
	10 36	5''	3×20^d	+19	1.14

^a X-F and X-L are exposures taken at the indicated position angle, for which the slit passes over RX J1856.5–3754 and star F or L (see Fig. 1); standards were taken at position angle 0°.

^b UT day in July 1999.

^c Parallactic angle at start of observation.

^d Third exposure taken with a GG 435 order separation filter.

^e Lost because of a glitch in telescope pointing.

on RX J1856.5–3754 appear similarly disappointing, with neither X-ray spectra taken with *Chandra* (Burwitz et al. 2001), nor ultra-violet spectra taken with *HST* (Pons et al. 2001) showing strong features.

The use of RX J1856.5–3754 to address the fundamental issues in physics and astrophysics mentioned above would benefit from – or even require – understanding the nature of the source. Walter et al. (1996) suggested it could be a young, cooling neutron star, or a neutron star kept hot by accretion from the interstellar medium. An alternative would be that it is a few million-year old magnetar, as was suggested for RX J0720.4–3125 on the basis of its 8.4-s periodicity (Kulkarni & van Kerkwijk 1998). An indication that RX J1856.5–3754 might be young is its proper motion, which Walter (2001) found to point away from the nearby Sco-Cen association. This has led him to the plausible suggestion that RX J1856.5–3754 was born in this association about a million years ago. For a young neutron star, however, it is hard to understand the lack of X-ray pulsations. Could it be that this neutron star has no significant magnetic field? Almost certainly, the interpretation of the high resolution X-ray spectra of this source will depend on knowing the composition of the atmosphere, the strength of the magnetic field, and the level of non-thermal emission.

In an effort to understand the nature of this important but enigmatic source, we have undertaken a series of observations. In this paper, we report on the first optical spectrum of RX J1856.5–3754 and on accurate optical-UV photometry. In our spectra, we find evidence for a nebula around RX J1856.5–3754. Those observations and the interpretation of the nebula will be the subject of the next paper (Van Kerkwijk & Kulkarni, in preparation).

Table 2. Log of the VLT imaging observations.

Object ^a	UT ^b day, time	Sequence ^c	Seeing (")	sec z
RX J1856	1, 06:57–07:52	BR α R α R	0.77	1.14–1.06
	07:54–08:48	BR α R α R	0.69	1.05–1.02
RX J1856	2, 07:23–08:16	BR α R α R	0.47	1.09–1.04
	08:19–09:11	BR α R α R	0.57	1.03–1.02
	09:14–09:45	BR α R	0.77 ^d	1.03–1.04
PG 0942	2, 23:10–23:20	BRBR	1.2	1.08
RX J1856	3, 06:00–06:54	BR α R α R	0.57	1.27–1.13
	06:59–07:51	BR α R α R	0.64	1.12–1.05
	08:00–08:52	BR α R α R	0.62	1.04–1.02
PG 1657	09:26–09:51	BRBR	1.3	1.60–1.69
	09:54–10:05	BRBR	1.0	1.71–1.81
RX J1856	4, 06:25–07:18	BR α R α R	0.75	1.19–1.09
	07:23–08:16	BR α R α R	0.75	1.08–1.03

^a In full: RX J1856.5–3754, PG 0942–029 and PG 1657+078.

^b UT day in May 2000.

^c For RX J1856.5–3754, the integration times were 138 s in B, 135 s in R, and 1019 s in H α ; for the standards, 3 and 20 s in B, and 1 and 20 s in R.

^d The B-band image has 1''.3 seeing.

The organisation of this paper is as follows. We describe our observations in Sects. 2 and 3, and the reduction in Sect. 4. We pay particular attention to accurate calibration, since some of our results turn out to be discrepant with previously published ground-based results. We also re-analyse the *HST* imaging, taking particular care to correct for systematic effects affecting faint stars. We discuss the results in Sects. 5 and 6.

2. Spectroscopic observations

We observed RX J1856.5–3754 on the night of 1999 July 15 to 16 at the 8-m Unit Telescope #1 (Antu) of the Very Large Telescope at Paranal, using the Focal Reducer/Low Dispersion Spectrograph FORS1 to obtain spectra covering the optical range and a number of images. A log of the observations is given in Table 1. The conditions were good, with the seeing varying from 0''.8 to 1''.2. The night started with patchy high cirrus, which disappeared later. The spectroscopy was done using a 300 lines mm^{−1} grism, the standard resolution collimator (final f-ratio of 3.13), and a Tektronix CCD detector with 2048 × 2048 pixels of 24 μ m. With this setup, the plate scale is 0''.2 pixel^{−1}, and the 3600–9000 Å wavelength range is covered at 2.6 Å pixel^{−1}. With the 1'' slit, the wavelength resolution is ~ 13 Å. The detector was read out through all four amplifiers, using the high gain setting ($\sim 1.6 \text{ e}^- \text{ ADU}^{-1}$) for the spectra, and the low gain setting ($\sim 3 \text{ e}^- \text{ ADU}^{-1}$) for the images.

Since the source is faint, we set up using two brighter stars nearby, stars F and L (here and below we follow the nomenclature of Walter et al. [1996]; see also Fig. 1), choosing the position angle such that the counterpart of RX J1856.5–3754 (hereafter star X) should fall in the slit as well. To measure the positions, we reanalysed the images

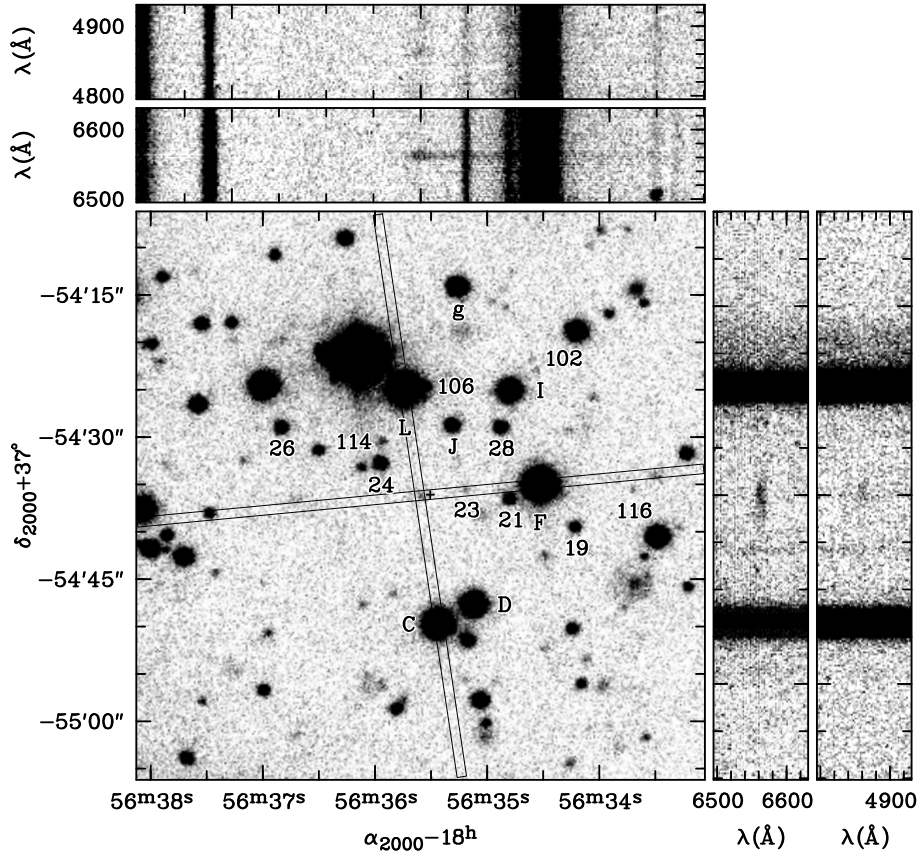


Fig. 1. Direct and spectral images of RX J1856.5–3754. The main panel shows the B-band image obtained at the start of our 1999 spectroscopic run (Table 1). The positions of the two settings of the slit for the spectroscopy are indicated by the two sets of parallel lines. The counterpart to RX J1856.5–3754 (star X) is within the intersection of the two slit positions. The position it had in October 1996, when it was observed with *HST* (Walter & Matthews 1997), is indicated by the plus sign. Previously identified stars for which we obtained photometry (see Table 5) are labeled below their image, except for stars C and 114 (label on the left) and stars D, I, 102, and 106 (label on the right). The panels above and to the right of the main panel show the part around H α and H β of the stacked spectral images for the two slit positions, over stars F and L, respectively. Along the spatial axes, the tick marks indicate 5'' intervals along the slit. The emission from the nebula around RX J1856.5–3754 can be seen clearly. The flux from RX J1856.5–3754 itself cannot be seen near H α .

of Walter & Matthews (1997), taken through F606W and F300W filters with the Wide Field Planetary Camera 2 (WFPC2) on board *HST*. The analysis of the *HST* images is described in more detail in Sect. 4.6.

In order to minimise slit losses due to differential refraction, we took set-up images through a Bessel B filter, used the differential refraction corrector of FORS1, and chose as reference the star for which the position angle was closest to the parallactic one. We obtained spectra with a total exposure time of 6.5 hour, taking individual exposures at various positions along the slit in order to mitigate the effect of bad pixels and other defects. To calibrate the slit losses, we followed our sequences of long integrations with short exposures through a wide slit (5'', formed using the multi-object slitlets). We observed two spectrophotometric flux standards through the wide slit to calibrate the instrumental response.

For verification of our set-up, we took a 5-minute B-band image before starting the first spectroscopic observation (which used star F as reference). From the image, star X appeared to have moved to the East.² After the first spectrum, therefore, we took three additional 5-minute B-band images³ to measure a more accurate position, and found that star X had, rather fortuitously, moved

exactly along the slit; hence, the first spectrum was not lost. For star F, we therefore kept the same set-up, while we changed the position angle for star L appropriately.

3. B, R, and H α imaging

Images in B, R, and H α of RX J1856.5–3754 and its environment were taken for us in the nights of 2000 April 30, May 1, May 2, and May 3, at Unit Telescope #2 (Kueyen) using the Focal Reducer/Low Dispersion Spectrograph FORS2. The standard resolution collimator was used, and a Tektronix CCD detector with 2048×2048 pixels of $24 \mu\text{m}$. The corresponding plate scale is $0''.2 \text{ pixel}^{-1}$.

During all nights the seeing was good, varying between $0''.55$ and $0''.8$. The nights of 2000 May 1 and May 2 were photometric, the night of April 30 mostly clear. Clouds appeared during the second half of the night of May 3, which includes the time used for RX J1856.5–3754; from the count rates, however, it appears that the clouds did not affect the observations. In total, ten series of images were taken of RX J1856.5–3754 (see Table 2). Nine were in the order B, R, H α , R, H α , R, and one – used to fill up time – was shortened to B, R, H α , R. Of the latter, the B-band image has much worse seeing and is not used. The total exposure time is about 21 minutes in B, 1 hour in R, and 5 hours in H α . During the photometric night, several standard fields of Landolt (1992) were observed. No useful separate H α calibration images were taken.

² In the meantime, the proper motion has been measured accurately using *HST* (Walter 2001).

³ Photometry of star X using these images was presented by Neuhäuser (2001). He finds $B = 25.1 \pm 0.4$, consistent with (though substantially less accurate than) our result.

4. Reduction

The data were reduced using the Munich Image Data Analysis System (MIDAS) and procedures running in the MIDAS environment. From bias frames taken before and after a given night, the bias level appeared somewhat variable, both in time and in position on the detector. However, relative to the levels found from the overscan pixels (determined separately for the four amplifiers), it remained constant. For bias subtraction, therefore, we subtracted both the levels from the overscan regions in individual frames and an average of the overscan-corrected bias frames (for the appropriate gain setting). The averages were determined separately for 1999 and 2000.

4.1. Photometry

All images were corrected for sensitivity variations using flat fields constructed from images of the sky taken at dusk and dawn (for the observations in 2000, only the dawn frames were used, since this produced much cleaner results). Averages were formed of the various series after filtering out cosmic-ray hits, verifying that even for the best-seeing images no stars were mistakenly affected.

Photometry was done by first determining the offset of instrumental magnitudes derived from the average frames from those derived from the images of the first sequence in the night of 2000 May 2, and then applying a calibration determined from the three standard fields observed during that night. We measured instrumental magnitudes using the DAOPHOT package (Stetson 1987). We used an iterative procedure, in which relatively isolated stars were selected and used to define a point-spread function (PSF), next the PSF was used to fit all stars and to subtract all but the PSF stars, and then the cleaned frame was used to determine an improved PSF, etc. We found that to model the variations in the PSF over the frame, a second-order dependence on position was required. Aperture corrections were determined from the difference between the fitted magnitudes and magnitudes measured in 20-pixel ($4''$) radius apertures on the PSF stars in the final frame in which all non-PSF stars had been removed.

The standard fields were analysed in two separate ways. For deriving the calibration using the Landolt (1992) photometry, we simply determined appropriate aperture magnitudes on the reduced frames (if not overexposed; in practice, we could only use the short frames). We inferred extinction coefficients of 0.126 and 0.076 in B and R, respectively, which are smaller than the typical values of 0.21 and 0.13 listed by ESO. We did not have sufficient data to measure the colour terms accurately, although we could confirm that the colour term for the B band is significant ($-0.025(B - R)$, i.e., the ESO B band is bluer than Landolt B), while the colour term for the R band is negligible. We estimate that the final uncertainty in the zero points is about 0.02 mag.

We also tried to calibrate our fields using fainter stars, since for many faint stars in Landolt fields,

Stetson (2000) has obtained calibrated magnitudes from archive observations. For this purpose, we analysed the frames using point-spread function fitting as described above. Unfortunately, however, while for the field of PG 1657+078, there are 32 stars with B-band magnitudes and 44 with R, for the field of PG 0942–029 there is only one star. As a result, we cannot derive an accurate solution including extinction terms, but only confirm the solution found using the Landolt photometry.

In Table 5, we list photometry for all point-like objects which are present in the *HST* Planetary Camera images taken through the F606W filter (Walter et al. 1996; Walter 2001), and which are detected in both B and R. One word of caution about the brightest stars, with $R \lesssim 18$ mag, which are overexposed in many of the R-band images. The magnitudes of these stars have been determined by PSF fitting to those pixels which were not overexposed, and are therefore somewhat more uncertain. We compared the magnitudes for these stars with magnitudes inferred from the first series of images, in which overexposure is less of an issue because of the relatively bad seeing. We found that the photometry in Table 5 may slightly underestimate the true brightness of the $R \lesssim 18$ mag stars, by ~ 0.01 mag.

4.2. Spectroscopy

For the spectroscopy, the flat fielding turned out to be problematic, because the spectroscopic flats, taken with the internal flat-field lamp illuminating the instrument cover, had an illumination pattern so different from that of the actual observations on sky that they were useless. Since the observations of the flux standards indicated that fringing was not a problem and that pixel-to-pixel sensitivity variations were much smaller than the sky-subtraction uncertainties for our very faint source, we decided to forego flat-fielding altogether. In order to equalise the four quadrants of the chip, which are read out through amplifiers with slightly different gain, we multiplied with the gains for the different amplifiers as measured by the instrument team. This provided very satisfactory results.

For the sky subtraction, clean regions along the slit within about 100 pixels of the sources of interest were selected, and these were fitted using a polynomial function. The order of the polynomial was mostly zero, but could be increased up to quadratic at any given column as long as further terms increased the goodness of the fit to the sky regions significantly. For each set of observations, the sky-subtracted images were registered and added together.

From both the individual and the summed sky-subtracted images, spectra were extracted using an optimal weighting scheme similar to that of Horne (1986). For this purpose, the spatial profile of the bright star (either F or L) was determined, and this was used to extract optimally weighted spectra at the position of the bright star itself, as well as at the position of star X. Furthermore, for verification, spectra were extracted at a number of empty positions. These were all consistent with zero flux.

The dispersion relation was found using an exposure of helium, argon and mercury lamps. Line positions were determined for positions along the whole slit. At a single position, a fourth order fit was sufficient, giving root-mean-square residuals of 0.5 \AA ; to obtain the same residuals for a two-dimensional relation required terms up to fifth order along the dispersion direction and second order along the spatial direction (for a total of 18 terms). The latter solution was used to calculate the wavelengths for all extracted spectra.

4.3. Spectral Images

In the extracted spectra of RX J1856.5–3754, emission lines of $H\alpha$ and $H\beta$ appeared. Inspection of the sky-subtracted frames showed that these lines were extended, especially along the slit over star F. In fact, it extended into the regions used to define the sky emission, and hence it had been partly removed in the sky-subtraction stage. In order to provide a cleaner picture, we rebinned the raw images to spectral images, in which every column is at a constant wavelength, removed cosmic rays, and formed averages for the two slit positions (excluding the first spectrum, which had a cosmic-ray hit at $H\beta$ near the target). Next, we determined the sky emission as a function of wavelength in regions far away from the neutron star, and subtracted this from all columns. The parts of the images around $H\alpha$ and $H\beta$ are shown in Fig. 1. From these spectral images, it is already clear that the neutron star has a nebula which is extended along the path it has travelled. This is confirmed by our $H\alpha$ imaging. The $H\alpha$ images and a discussion of the nature of this nebula will be presented elsewhere (Van Kerkwijk & Kulkarni, in preparation).

4.4. Flux calibration of the spectra

For the flux calibration, the spectra were first corrected for atmospheric extinction using the average La Silla extinction curve. While this will be only approximately correct, it facilitates the next step, the determination of the slit losses. For this purpose, the ratio with the wide-slit spectra was formed for each of the bright star spectra taken through the $1''$ slit. These ratio spectra were approximated with second-degree polynomials, which were used to correct all spectra.

Finally, the spectra were corrected for the response of the spectrograph derived from the observation of the spectrophotometric standard EG 274 (Hamuy et al. 1992, 1994). The spectra of this standard were extracted in the same manner as described above, but in addition a correction was made for the blue second-order light that overlaps the part of the spectrum at $\gtrsim 6500 \text{ \AA}$ (the correction was determined with the help of the spectrum taken through the GG 435 filter). We observed the DA white dwarf BPM 16274 as an additional calibrator. Unfortunately, we realised later that this star is only calibrated in the ultraviolet. We still used it to verify our response curve us-

ing a model spectrum kindly provided by D. Koester (for $T_{\text{eff}} = 18750 \text{ K}$ and $\log g = 7.83$, as inferred by Bragaglia et al. [1995], and normalised to $V = 14.20$, as measured by Eggen [1969]). For wavelengths longer than $\sim 3700 \text{ \AA}$, the comparison was very satisfactory, as was a similar comparison using a DA white-dwarf model provided by D. Koester for EG 274 ($T_{\text{eff}} = 24250 \text{ K}$, $\log g = 8.05$; Vauclair et al. 1997).

The above gives us confidence that the relative calibration over the spectral range is accurate. The accuracy of the absolute calibration, however, is less clear, since the spectra of EG 274 were taken in the beginning of the night, when some patchy cirrus was still present. In order to assess the influence of the cirrus, we compared fluxes from all B-band (acquisition) images and all wide-slit spectra. We found that there were variations before about 1 UT, but that after that time the measured count rates indicate the sky was clear. To see whether our flux calibration was influenced by the cirrus, we used the B, V, and R filter curves of Bessell (1990) to determine synthetic B, V, and R-band magnitudes for all brighter objects in our wide-slit spectra. For BPM 16274, we find $V = 14.20$, $B - V = -0.04$, quite consistent with the observed $V = 14.20$, $B - V = -0.015$ (Eggen 1969). Also for stars F (twice), L, and C (in the slit for the star L position; see Fig. 1), the synthetic magnitudes are in good agreement with our B and R-band photometry, as can be seen in Table 3. We conclude that the absolute calibration of our spectra is accurate to 0.02 mag.

4.5. Previous photometry of stars L, C, and F

While our spectrophotometry of stars L, C, and F agrees well with our own photometry, it disagrees with measurements in the literature: the synthetic V-band magnitudes are 0.4 mag brighter than both the V-band magnitudes of Neuhäuser et al. (1997) and the V-band magnitudes inferred from Gunn g and r measurements of Campana et al. (1997). (The magnitudes of Walter et al. (1996) differ even more, by ~ 1 mag, but Neuhäuser et al. have already noted that Walter et al. used an incorrect zero point.)

Comparing colours, we find that our synthetic $B - V$ and $V - R$ colours are systematically redder and bluer, respectively, than those of Neuhäuser et al. (1997). The inferred $B - R$ values, however, are consistent. This suggests there may be a problem with the V-band only. Indeed, the V, $B - V$, and $V - R$ values listed by Neuhäuser et al. imply B and R-band magnitudes that are in reasonable agreement with our values for bright stars like stars L, C, and F. If our synthetic B and R are correct, however, our synthetic V should be correct too, since the three bands are tied to each other by *relative* calibration on wide-slit spectra, which proved very reliable on BPM 16274.

In order to settle the issue, we classified the spectra (see Fig. 2), using the spectral atlases of Silva & Cornell (1992) and Torres-Dodgen & Weaver (1993). All three stars appear to be of spectral type G; see Table 3. For all three, our colours are consistent with the spectral types

Table 3. Spectral types, colours and magnitudes for stars L, C, and F. For each star, the first line lists the spectral type inferred from the spectra and the corresponding intrinsic colours (for a main-sequence star; Cox 2000). The following lines list synthetic magnitudes and colours derived from the wide-slit spectra (two for star F), and magnitudes derived from our images. All magnitudes and colours have ~ 0.02 mag uncertainty.

Star	Spectral type	B	$(B - V)$	V	$(V - R)$	R
L	G3–6 V–III	0.64–0.66		0.36–0.37		
		18.07	0.74	17.33	0.41	16.92
		18.06				16.93
C	G6–8 V–III	0.66–0.76		0.37–0.44		
		18.56	0.78	17.78	0.43	17.35
		18.53				17.34
F	G9–K0 V–III	0.79–0.82		0.45–0.47		
		17.94	0.88	17.06	0.46	16.60
		17.93	0.86	17.07	0.46	16.61
		17.91				16.58

(for a small amount of reddening; see Sect. 5), while the $B - V$ and $V - R$ (but not $B - R$) colours of Neuhäuser et al. (1997) and the $g - r$ colour of Campana et al. (1997) are inconsistent with the spectral types (independent of reddening). We thus conclude that, despite the inconsistencies with earlier work, our calibration is reliable.

4.6. Archival *HST* imaging

RX J1856.5–3754 has been a regular target for WFPC2 observations with *HST*, both to measure broad-band photometry (Walter & Matthews 1997; Pons et al. 2001) and to determine the proper motion and parallax (Walter 2001). We reanalysed all images (see Table 4), both to provide a final verification of our flux calibration, and to extend the optical spectral energy distribution for RX J1856.5–3754 to shorter wavelengths. In our analysis, we take into account that during the last years updated zero points have become available (Baggett et al. 1997), and that new prescriptions have been published for correcting for changes in the amount of contaminants on the CCD windows (Baggett & Gonzaga 1998), for the so-called “long-versus-short anomaly” (Casertano & Mutchler 1998), and for the slowly degrading charge-transfer efficiency (CTE; Whitmore et al. 1999); the latter two are particularly important for faint objects. Furthermore, recently a package specifically written for WFPC2 photometry, HSTPHOT, has been made available by Dolphin (2000a), which includes many of the above corrections (Dolphin 2000b).

Our analysis started with the pipe-line reduced WFPC2 images. We measured photometry using HSTPHOT, as well as, for comparison, our own procedures. For the HSTPHOT reduction, we followed the prescription of Dolphin (2000a): (i) mask bad pixels; (ii) combine images taken at the same position and remove

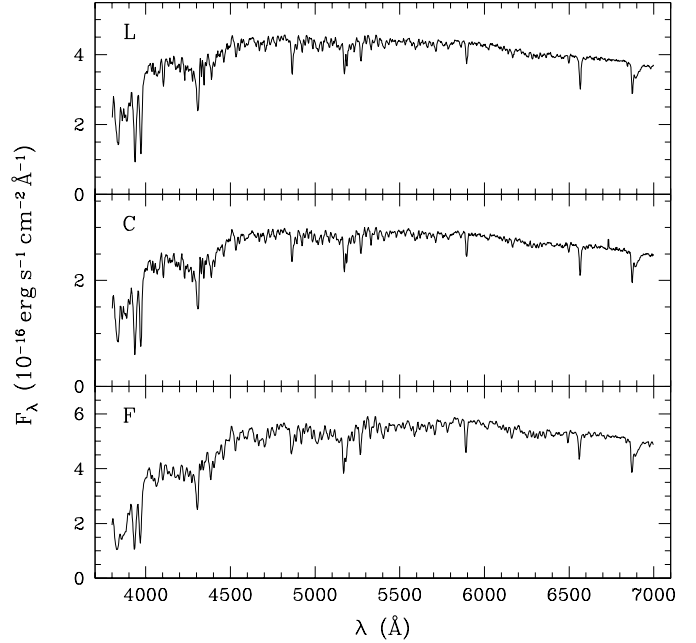


Fig. 2. Spectra of stars L, C, and F. These are used for classification and verification of the flux calibration in Sect. 4.5; see also Table 3.

cosmic ray hits; (iii) determine sky levels; (iv) remove hot pixels; and (v) measure photometry by point-spread function fitting on the combined image(s). In the last step, we disabled the determination of point-spread function residuals and aperture corrections for the F170W image exposures, since these lack a sufficient number of well-exposed stars. For the F170W exposures, we also had to make a change to the source code for the sky determination, viz., to remove the constraint that the fitted sky level had to be positive. While this constraint is physically reasonable, slight inadequacies in the pipe-line subtraction of the bias and dark current can lead to negative count rates, which, if not corrected for, lead one to underestimate a source’s brightness; this is indeed the case for the F170W images. Another change we made to the source code was that we forced the use of the published (Dolphin 2000b) charge-transfer efficiency corrections for all bands (the MULTIPHOT routine in the HSTPHOT distribution uses more recent corrections for the optical filters, but not for the ultraviolet ones; the difference is rather small).

In Table 5, we list the results. For the F300W and F606W filters, the averages of the individual measurements are listed (no significant variability was found for any star). Comparing the *HST* photometry with our ground-based results, we find that the two are consistent.

To verify our technique, we also measured aperture photometry on averaged images, produced in the manner described by Van Kerkwijk et al. (2000). We measured count rates in apertures with a range of radii, derived aperture corrections to the standard $0''.5$ radius aperture, calculated the corrections discussed above (following the

Table 4. Log of archival *HST*/WFPC2 imaging. The F606W and F450W images are taken at two dither positions, different by 5.5 pixels in both X and Y, while the F300W and F170W images are all taken at the same position

Identifier	UT	Filter	t_{int}
u3im010[1–4]	1996 Oct 6	F606W	$3 \times 1000 + 1400$
u3im010[5–6]		F300W	2×1200
u51g010[1–8]	1999 Mar 30	F606W	8×900^a
u51g030[1–4]		F170W	$2 \times 2800 + 2 \times 2600$
u51g040[1–8]	26	F450W	8×900
u51g020[1–2]	Sep 16	F300W	2×1300
u51g020[3–6]		F606W	$3 \times 1300 + 1290.5$

^a u51g0106 has a high background and a large number of cosmic-ray hits. It has not been used in the analysis.

above references and the *HST* data and WFPC2 instrument handbooks), and converted to calibrated magnitudes by applying a 0.1 mag aperture correction from the 0.5 aperture to infinity and using the zero points of Baggett et al. (1997). We note that, as alluded to above, some corrections are large, especially for faint objects. For star X, the “long-vs-short” corrections are -0.17 , -0.3 , -0.3 , and -0.4 mag in F606W, F450W, F300W, and F170W, respectively (but see below); the 1996 CTE corrections are -0.07 and -0.13 mag in F606W and F300W, respectively; and the 1999 CTE corrections are -0.11 , -0.14 , -0.30 , and -0.23 mag in F606W, F450W, F300W, and F170W, respectively.

Comparing the results with those derived using HSTPHOT, we found good consistency for the brighter stars. For the fainter stars like star X, however, this was only the case if we did not apply the “long-versus-short” correction. Indeed, no such correction is applied in HSTPHOT, since Dolphin (2000b) has not found any evidence for it; he argues that its appearance likely reflects inaccurate sky subtraction in the procedures used by Casertano & Mutchler (1998). We do not have strong independent evidence either way, but note that if we do apply the correction, our VLT photometry for faint objects like star X becomes inconsistent with the *HST* results (while the results remain consistent for the brighter stars, since these are not affected).

Comparing to the magnitudes listed by Walter & Matthews (1997) and Walter (2001), we find that our results are roughly consistent for the brighter stars, but that they differ for the fainter stars; in particular, for star X, while the F606W magnitude is virtually identical, our F300W magnitude is 0.4 mag brighter than that of Walter & Matthews (note that these authors list magnitudes on the ST system). We suspect that the differences largely reflect our use of up-to-date corrections. This suspicion is strengthened by the fact that Pons et al. (2001), after making similar corrections, also find fluxes that differ from those of Walter & Matthews (1997). Indeed, their results are very similar to ours (their fluxes are fainter

Table 5. Photometry of RX J1856.5–3754 and stars in the field. All magnitudes are in the Vega system. For star X, the uncertainties are as listed. For all other stars, an indication of the uncertainty is given by the presence of no ($\sigma < 0.035$ mag), one ($0.035 < \sigma < 0.075$) and two colons ($0.075 < \sigma < 0.125$). The zero point uncertainty is about 0.05 mag for the F170W band, and 0.02 mag for all others.

Object ^a	m_{170}	m_{300}	B	m_{450}	m_{606}	R
X	21.37 ± 0.10	23.00 ± 0.07	25.19 ± 0.07	25.21 ± 0.06	25.64 ± 0.02	25.80 ± 0.11
C	21.62::	19.30	18.53	18.43	...	17.34
D	...	20.80	19.28	19.07	...	17.69
F	20.70::	19.01	17.91	16.58
I	...	21.12	19.60	19.49	...	17.89
J	...	23.42::	21.83	21.64	20.39	20.08
L	20.93::	18.72	18.06	16.93
g	...	20.74	20.48	20.44	19.72	19.48
19	22.95	22.66	21.40	20.99
21	22.60	22.25	20.48	19.97
23	25.56::	24.97:	23.26	22.68
24	21.98	21.62	19.80	19.34
26	22.28	22.19	20.90	20.56
28	...	23.10:	22.16	21.93	21.05	20.79
102	25.77::	25.37::	23.93	23.49
106	23.13	22.84	21.83	21.53
114	24.66:	24.88:	23.88	23.31
116	25.35::	...	23.55	23.05
Δ_{ST}^b	-0.412	-0.024	-0.602	-0.475	+0.316	+0.555
$\langle \lambda \rangle^c$	1760	2955	4371	4520	5933	6464
$\langle A_\lambda / A_V \rangle^c$	2.648	1.931	1.336	1.297	0.949	0.841

^a X is the optical counterpart. For the other objects, the nomenclature follows Walter et al. (1996, uppercase letters), Campana et al. (1997, two digit numbers), Walter (2001, three-digit numbers) and ourselves (object g). See Fig. 1.

^b Magnitude differences between the ST and Vega systems. These can be used to calculate the flux by $f_\lambda = 10^{-0.4(m + \Delta_{\text{ST}} + 21.100)} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$. From Baggett et al. (1997, WFPC2 filters) and Bessell et al. (1998, B and R).

^c Effective wavelength and reddening for a Rayleigh-Jeans ($f_\lambda \propto \lambda^{-4}$) spectrum. See App. A.

by 0.04...0.10 mag and their uncertainties slightly larger, since they use aperture photometry rather than point-spread function fitting).

5. Spectrum and spectral energy distribution

The reduced, flux-calibrated spectrum is shown in Fig. 3. We limit ourselves to the wavelength range of 3800–7000 Å, since at shorter wavelengths the flux calibration becomes unreliable, while at longer wavelengths second-order light starts to contribute significantly for objects as blue as star X (it is negligible shortward of 6500 Å and rises approximately linearly to 10% at 7000 Å). We also removed 20 Å-wide regions around H α , H β , and H γ , which are (or might be in case of H γ) contaminated by nebular emission.

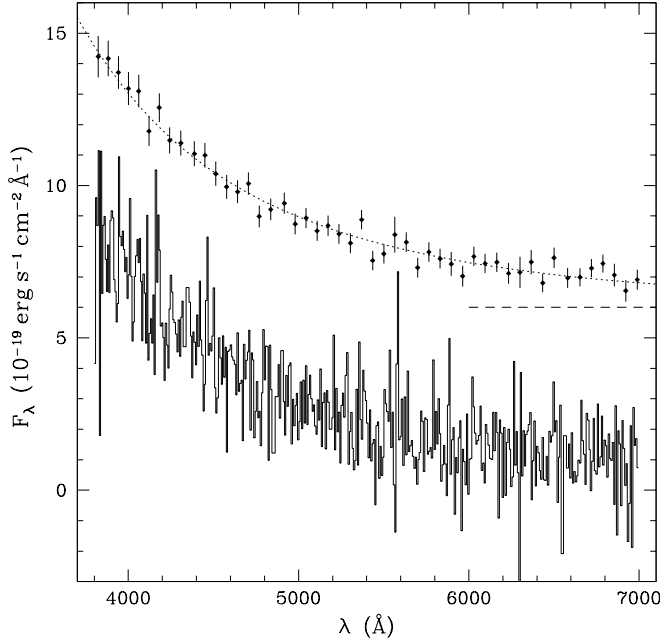


Fig. 3. Optical spectrum of RX J1856.5–3754. For the histogram, the data were averaged in 3-pixel wide bins ($\sim 7.5 \text{ \AA}$ or $\sim 60\%$ of a resolution element). The points with (Poisson) error bars reflect averages over 25-pixel ($\sim 63 \text{ \AA}$) wide bins. These are offset vertically; their zero level is indicated by the dashed line. The dotted line indicates the absorbed Rayleigh-Jeans spectrum that best fits the photometry. The observed spectrum is consistent with this and has no significant features.

The spectrum does not contain any significant features. The best limits to the equivalent width of any feature are obtained shortward of $\sim 5400 \text{ \AA}$: about 16 \AA for features with 50% depth beneath the continuum, and about 60 \AA for features with 25% depth.

The spectrum is shown together with the photometry in Fig. 4. One sees that the spectrum is consistent with both the VLT and *HST* photometry. Indeed, integrating the spectrum over the B and F450W bandpasses, we infer $B = 25.22$ and $m_{450} = 25.25$, which compares well with our photometry (Table 5), giving additional confidence in the calibration of all three data sets.

Both spectrum and photometry indicate a spectral energy distribution close to that of a Rayleigh-Jeans tail, as would be expected for this very hot source. Assuming an intrinsic Rayleigh-Jeans spectrum, we determine the reddening to the source by fitting a reddened λ^{-4} spectral distribution to the photometry. We do not include the spectrum in the fit, because we consider its absolute flux calibration somewhat less reliable and also because we wish to use it to verify the result from the photometry. We use the following relation between Vega magnitude m_{Vega} and input spectrum f_{λ} ,

$$m_{\text{Vega}} = -2.5 \log \left(\frac{\int f_{\lambda} \frac{\lambda}{hc} R_{\lambda} d\lambda}{\int \frac{\lambda}{hc} R_{\lambda} d\lambda} \right) - 21.100 - \Delta_{\text{ST}}, \quad (1)$$

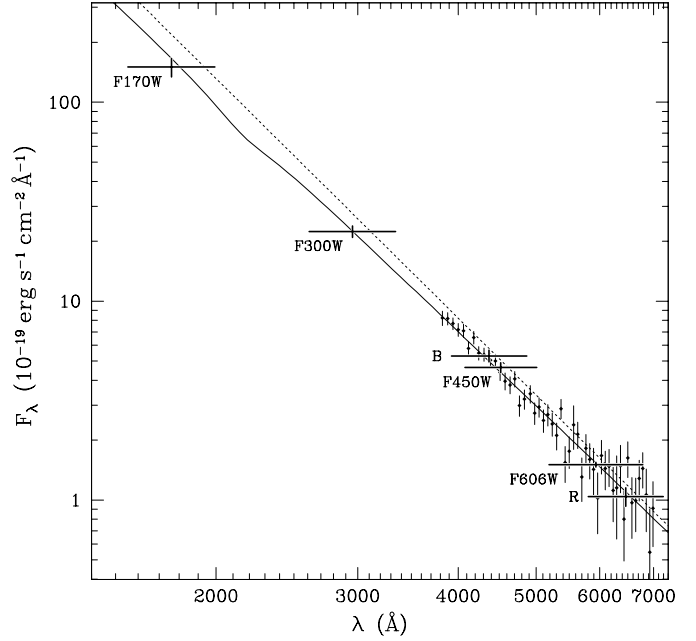


Fig. 4. Optical/ultraviolet spectral energy distribution of RX J1856.5–3754. The thick-set points indicate fluxes derived from VLT and *HST* photometry. The vertical lines indicate the 1σ errors while the horizontal ones are measures of the filter widths. Overdrawn are the best-fit absorbed (drawn curve) and unabsorbed (dotted curve) Rayleigh-Jeans spectra, as well as the spectrum observed with the VLT, averaged in 25-pixel wide bins ($\sim 63 \text{ \AA}$).

where R_{λ} is the response at wavelength λ , and Δ_{ST} is the magnitude difference between the Vega and ST systems (see Table 5). In the latter system, a flat spectrum has the same magnitude in all filters; it matches the Johnson system at V. The integrations are done over photon rates, since what we measure with CCDs is the rate of photons in a particular band. The response curves we used for the WFPC2 filters are the system response curves for the planetary camera, taken from the Space Telescope Science Institute web site. For the B and R filters, we used the Landolt filter curves from Bessell (1990). For the input spectrum, we take

$$f_{\lambda} = f_{\lambda_0} \left(\frac{\lambda}{\lambda_0} \right)^{-4} 10^{-0.4(A_{\lambda} - A_{\lambda_0})}, \quad (2)$$

where f_{λ_0} is the observed flux at reference wavelength λ_0 and A_{λ} is the reddening. We use the reddening curve of Cardelli et al. (1989) for $R = 3.1$; we include corrections for the optical as described by O'Donnell (1994). To minimise the covariance between f_{λ_0} and A_V , we chose $\lambda_0 = 5000 \text{ \AA}$.

From the fit to the photometry, we find $f_{5000} = (2.96 \pm 0.06) \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$ and $A_V = 0.12 \pm 0.05$; the fit is acceptable, with $\chi^2_{\text{red}} = 3.5$ for four degrees of freedom (six bands and two parameters; note that for the uncertainties we used the measurement errors with the zero-point uncertainties added in quadrature; see Table 5).

For the reddening curve used, $A_{5000}/A_V = 1.138$, and thus the unabsorbed flux is $f_{\lambda_0,0} = (3.36 \pm 0.17) \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$, where the error is dominated by the uncertainty in A_V . The fit is shown in Fig. 4; it can be seen that it also is a good fit to the optical spectrum (see also Fig. 3), with $\chi^2 = 37$ for 49 degrees of freedom (no free parameters; also for other choices of binning, one finds $\chi_{\text{red}}^2 \simeq 1$).

The inferred reddening is consistent with the range $A_V \simeq 0.05 \dots 0.12$ expected⁴ from the range in X-ray column density $N_{\text{H}} = (1.0 \dots 2.2) \times 10^{20} \text{ cm}^{-2}$ found from different model fits to the X-ray and EUVE spectrum (Pons et al. 2001; Burwitz et al. 2001). It is also consistent with the limit set by the total amount of reddening along this line of sight, which we can infer from stars L, C, and F. From the difference between the observed colours and the intrinsic colours for these stars (inferred from their spectral types; see Table 3), we infer $E_{B-R} = 0.10 \pm 0.03$, which corresponds to $A_V = 0.20 \pm 0.06$. Using this reddening, and assuming L, C, and F are main-sequence stars with absolute magnitudes $M_V \simeq 5.0, 5.4$, and 5.8 (Cox 2000), respectively, their distances are 2.7, 2.8, and 1.6 kpc, respectively⁵. Thus, they are well in the background relative to star X, as well as relative to the CrA cloud complex, which, apparently, contributes very little extinction in this line of sight, unlike what was suggested previously (Walter et al. 1996).

6. Deviations from a Rayleigh-Jeans tail?

Above, we assumed a λ^{-4} spectral energy distribution for star X. Indeed, from the observations, we cannot determine the slope of the spectrum independently: spectra with $f_{\lambda} = f_{\lambda_0} (\lambda/\lambda_0)^{-\alpha}$ reproduce the photometry equally well as long as $A_V = 0.12 \pm 0.06 + 0.89(\alpha - 4)$. However, stars L, C, and F pose an upper limit of 0.20 mag on the reddening to star X. Using $0 < A_V < 0.20$, one infers that $3.8 \lesssim \alpha \lesssim 4.1$, i.e., the spectrum has to have a slope very close to that of a Rayleigh-Jeans tail.

6.1. Limits to a non-thermal component

Our data place stringent constraints on any contribution from non-thermal emission. For instance, fitting a sum of a Rayleigh-Jeans tail and a non-thermal spectrum with $f_{\nu} \propto \nu^{0.1}$ (i.e., $f_{\lambda} \propto \lambda^{-2.1}$), such as observed for the Crab pulsar (Sollerman et al. 2000; Carramiñana et al. 2000),

⁴ We used $A_V = 5.6(N_{\text{H}}/10^{22} \text{ cm}^{-2})$, as inferred by Predehl & Schmitt (1995) for distant, highly absorbed objects. In these, the line of sight samples all phases of the interstellar medium; hence, the relation may not be valid locally.

⁵ If they are giants, with $M_V \simeq 0.8$ (Cox 2000), they would be at ~ 20 kpc and might be associated with the Sagittarius dwarf galaxy (Ibata et al. 1995), which is at ~ 25 kpc and $\sim 7^\circ$ away from RX J1856.5–3754. Indeed, stars L and C have radial velocities of $+90 \pm 70 \text{ km s}^{-1}$, consistent with being associated. Star F, with a radial velocity of $-130 \pm 70 \text{ km s}^{-1}$, almost certainly is not associated.

the best-fit has zero contribution from non-thermal emission and the same reddening as inferred above. Even for zero reddening, the best-fit contribution is only 11% at 6000 Å (for this fit, $\chi^2 = 5$). The 95% confidence upper limit to a non-thermal contribution for a Crab-like spectrum is 20% at 6000 Å (for a non-thermal spectrum with $f_{\nu} \propto \nu^{-0.8}$, as observed for Geminga [Martin et al. 1998], this reduces to 8%).

6.2. Limits to emission from an accretion disk

We also considered whether an accretion disk might be present, from which RX J1856.5–3754 is accreting. Accretion from “debris disks” has been invoked in models of anomalous X-ray pulsars (e.g., Van Paradijs et al. 1995; Chatterjee et al. 2000). Furthermore, Perna et al. (2000) suggested that the deviation from a Rayleigh-Jeans spectrum found from optical observations of PSR B0656+14 (Koptsevich et al. 2001 and references therein) could be due to the presence of such a disk. For RX J1856.5–3754, we considered two cases. For the first, we assumed the source is powered by accretion from a disk, in which case both viscous heating (Shakura & Sunyaev 1973) and irradiation by the neutron star (Vrtilek et al. 1990) lead to optical emission. We used routines described by Hulleman et al. (2000a,b) to calculate the emission from both processes⁶, integrating between an inner radius r_{in} and an outer radius of 10^{14} cm . For the neutron star spectrum, we take the black-body fit of Pons et al. (2001) that best fits the observed X-ray to optical spectral energy distribution ($kT = 48 \text{ eV}$, $R/d = 0.11 \text{ km/pc}$) and a distance of 60 pc as inferred from the parallax (Walter 2001). We found that for a disk extending all the way in to the neutron star ($r_{\text{in}} = 10^6 \text{ cm}$), the optical emission predicted far exceeds that observed, by three orders of magnitude in R. In order for the emission to remain below 10% of the R-band flux (i.e., $R_{\text{disk}} \gtrsim 28 \text{ mag}$), the inner radius had to be $\gtrsim 2 \times 10^9 \text{ cm}$. This could be the radius where the disk is disrupted by a magnetic field; if so, and if the neutron star were rotating at equilibrium, its period would have to be $\gtrsim 40 \text{ s}$. If such a disk were present, its spectrum would be very red. At the limit, one would predict $J = 22$ and $K = 19 \text{ mag}$.

In principle, the neutron star could have a disk even if the X-ray emission is not due to accretion, in which case the mass accretion rate could be lower and hence the viscous heating of the disk much reduced. In the second case we considered, therefore, we ignored the contribution of viscous heating. Also for this case, the disk cannot extend all the way in to the neutron star (it would still exceed the observed R-band flux by an order of magnitude); we find

⁶ Assuming that the disk is optically thick and radiates as a black body. It is questionable whether these assumptions are justified for the very low accretion rate considered here, but given the speculative nature of the presence of an accretion disk, we felt a more detailed analysis was not warranted.

$r_{\text{in}} \gtrsim 3 \times 10^8 \text{ cm}$ ($\Rightarrow P \gtrsim 3 \text{ s}$; at the limit, the predicted infrared magnitudes are $J = 23$ and $K = 20$).

We conclude that all measurements are consistent with a slightly reddened Rayleigh-Jeans spectrum, with no evidence for features or for a contribution from non-thermal emission or an accretion disk.

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Appendix A: Effective wavelengths and reddening coefficients

To ease future fitting (and plotting), we list in Table 5 the wavelengths at which one should evaluate the flux for spectra which are close to Rayleigh-Jeans, as well as relative reddening coefficients. These were calculated as follows

$$\langle \lambda \rangle = \left(\frac{\int \lambda^{-4} \lambda R_{\lambda} d\lambda}{\int \lambda R_{\lambda} d\lambda} \right)^{-1/4}, \quad (\text{A.1})$$

$$\left\langle \frac{A_{\lambda}}{A_V} \right\rangle = \frac{-2.5}{0.12} \log \frac{\int 10^{-0.4 \times 0.12 (A_{\lambda}/A_V)} \lambda^{-4} \lambda R_{\lambda} d\lambda}{\int \lambda^{-4} \lambda R_{\lambda} d\lambda} \quad (\text{A.2})$$

Note that the wavelength $\langle \lambda \rangle$ is not the effective wavelength of the detected photons in a given filter, but is the wavelength at which a λ^{-4} spectrum has the same flux as a λ^0 spectrum that produces the same number of counts in the filter considered.

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